XIV Serbian Conference on Spectral Line Shapes in Astrophysics Bajina Bašta, Serbia, June 19 - 23, 2023 Book of Abstracts, Eds. Luka Č. Popović, Nataša Bon, Edi Bon and Sylvie Sahal-Bréchot

## INITIAL STARK WIDTH INVESTIGATION OF TeII SPECTRAL LINES AND THEIR IMPORTANCE IN ASTROPHYSICAL APPLICATIONS

Z. Majlinger<sup>1,2,3</sup>, M. S. Dimitrijević<sup>1</sup>, V. A. Srećković<sup>4</sup>

<sup>1</sup>Astronomical Observatory, Volgina 7, 11060 Belgrade, Serbia
<sup>2</sup>ES "I.G. Kovačić", Kralja Tomislava 18, 51326 Vrbovsko, Croatia
<sup>3</sup>ES "Vladimir Gortan", Prilaz Vladimira Gortana 2, 51000 Rijeka, Croatia
<sup>4</sup>Institute of Physics, UB, P.O. Box 57, Belgrade, Serbia

E-mail: zlatko.majlinger@gmail.com, vlada@ipb.ac.rs, mdimitrijevic@aob.rs

Set of 39 spectral lines of TeII have been prepared for a purpose of pressure broadening parameter determination. For some of these lines, Stark widths were calculated using modified semiempirical (MSE) approach (Dimitrijević and Konjević, 1980). Line and multiplet factors were taken from Shore and Menzel (1965), while matrix elements were determined by using of oscillator strengths (Zhang et al, 2013) whenever it was possible similarly as it is done, for example, in the analysis of Stark broadening of Lu III (Majlinger et al, 2015). These results will be compared with the results obtained by classical use of Bates-Damgaard method within MSE formalism (Bates and Damgaard, 1949). Stark widths of Te II can be helpful in analysis of astrophysical spectra, where MSE method has been already used successfully in the previous researches concerning the Stark broadening calculations of spectral lines.

## References

- [1] Bates, D.R., Damgaard, A., 1949, Phil. Trans. Roy. Soc. London, 242A
- [2] Dimitrijević, M. S., Konjević, N., 1980, J. Quant. Spect. Radiat. Transfer, 24, 454
- [3] Shore, B.W., Menzel, D. H., 1965, Astrophys. J. Suppl. Series, 12, 187
- [4] Majlinger, Z., Simić, Z., Astron. Astrophys, 2015, 36, 671
- [5] Zhang W., Palmeri P., Quinet P., 2013, Astron. Astrophys. 551A, 136